

THE MOST ANCIENT SPIRAL GALAXY: A 2.6-GYR-OLD DISK WITH A TRANQUIL VELOCITY FIELD

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ApJ accepted 2017 October 18

ABSTRACT

We report an integral-field spectroscopic (IFS) observation of a gravitationally lensed spiral galaxy *A1689B11* at redshift $z = 2.54$. It is the most ancient spiral galaxy discovered to date and the second kinematically confirmed spiral at $z \gtrsim 2$. Thanks to gravitational lensing, this is also by far the deepest IFS observation with the highest spatial resolution (~ 400 pc) on a spiral galaxy at a cosmic time when the Hubble sequence is about to emerge. After correcting for a lensing magnification of 7.2 ± 0.8 , this primitive spiral disk has an intrinsic star formation rate of $22 \pm 2 M_{\odot} \text{ yr}^{-1}$, a stellar mass of $10^{9.8 \pm 0.3} M_{\odot}$ and a half-light radius of $r_{1/2} = 2.6 \pm 0.7$ kpc, typical of a main-sequence star-forming (SF) galaxy at $z \sim 2$. However, the $H\alpha$ kinematics show a surprisingly tranquil velocity field with an ordered rotation ($V_c = 200 \pm 12 \text{ km s}^{-1}$) and uniformly small velocity dispersions ($V_{\sigma, \text{mean}} = 23 \pm 4 \text{ km s}^{-1}$ and $V_{\sigma, \text{outer-disk}} = 15 \pm 2 \text{ km s}^{-1}$). The low gas velocity dispersion is similar to local spiral galaxies and is consistent with the classic density wave theory where spiral arms form in dynamically cold and thin disks. We speculate that *A1689B11* belongs to a population of rare spiral galaxies at $z \gtrsim 2$ that mark the formation epoch of thin disks. Future observations with JWST will greatly increase the sample of these rare galaxies and unveil the earliest onset of spiral arms.

Subject headings: cosmology: observations — galaxies: formation — galaxies: evolution — galaxies: high-redshift — galaxies: spiral

1. INTRODUCTION

One of the most common features of disk galaxies in the local universe is the presence of prominent spiral arms. Among millions of galaxies charted in the local universe, $\sim 70\%$ exhibit spiral arms (e.g., Nair & Abraham 2010; Willett et al. 2013). However, the number density of spiral galaxies decreases dramatically at high redshift (Conselice 2014; Elmegreen & Elmegreen 2006). For example, only one spiral galaxy has been spectroscopically confirmed at $z \gtrsim 2$ (Law et al. 2012).

Spiral arms serve important purposes in galaxy formation and evolution: they are sites of star formation and are intimately associated with the formation of the thin and thick disk (Elmegreen 2011; Conselice 2014; Martinez-Medina et al. 2015). Spiral arms play an active role in driving the radial and azimuthal mixing of the metals, redistributing angular momentum, and smoothing out small-scale mass distributions (e.g., Sellwood & Binney 2002; Sellwood 2014; Grand et al. 2015, 2016). The number and pitch angle of spiral arms are strongly correlated with the mass distribution of the disk and can be a powerful tool to constrain the bulge and black hole masses (Athanasoula et al. 1987; Kennicutt 1981; Elmegreen & Elmegreen 1990; Berrier et al. 2013; Dobbs & Baba 2014; Seigar et al. 2014; Davis et al. 2015, 2017). The onset of spiral structures offers crucial insights into the origin of the Hubble sequence (Driver et al. 1998;

Cen 2014; Genel et al. 2015).

The necessary and sufficient conditions for spiral arm formation remain inconclusive, despite major developments in the 1960s and decades of studies (e.g., Toomre 1977; Athanasoula 1984; Sellwood 2011; Dobbs & Baba 2014). Popular mechanisms for spiral arm formation are largely based on early analytical works: e.g., the density wave theory (Lindblad 1960; Lin & Shu 1964; Kalnajs 1971), swing amplifications (Goldreich & Lynden-Bell 1965; Julian & Toomre 1966), and bars and tidal interactions (Kormendy & Norman 1979; Salo & Laurikainen 1993). These three mechanisms are not mutually exclusive and have mixed observational successes (e.g., Sellwood 2011; D’Onghia et al. 2013; Shu 2016; Pour-Imani et al. 2016). Whether spiral arms are long-lived patterns or transient features is still hotly debated in theory and poorly constrained in observations (Sellwood & Binney 2002; Sellwood 2011). The progress in the theory of spiral arm formation is slow and current efforts primarily focus on nearby galaxies (Dobbs & Baba 2014).

Breakthroughs can come from observations of high-redshift galaxies when spiral arms are in the early stages of formation. All classic analytical models of spiral arm formation assume an infinitesimally thin and cold disk in a stable rotation (e.g., Toomre 1977; Bertin & Arnouts 1996; Rafikov 2001; Sellwood 2014). It is unclear if this assumption holds at high redshift. Compared to local galaxies, high-redshift disks are gas-rich, globally unstable, and tend to have larger velocity dispersions and thicker disks (e.g., Freeman & Bland-Hawthorn 2002;

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Elmegreen & Elmegreen 2006; Law et al. 2007; Tacconi et al. 2013; Glazebrook 2013; Wisnioski et al. 2015; Johnson et al. 2017; Zhou et al. 2017). The role of the interstellar medium (ISM) and gas feedback in shaping spiral arms is usually over-simplified in local galaxies, but could become increasingly important and complicated at high redshift (e.g., Bertin & Romeo 1988; Elmegreen & Thomasson 1993; Wada et al. 2011; Ghosh & Jog 2015). Spiral arms are also sensitive to external processes such as bars, galaxy mergers, and gas accretion, all of which are different at high redshift (Sellwood 2004; Wada et al. 2011; Martig et al. 2012; Dobbs & Baba 2014). Studying spiral galaxies at a dynamically hostile cosmic time has the unique advantage of probing the most sensitive factor(s) responsible for spiral arm formation.

Spiral galaxies are rarely observed at $z \gtrsim 2$ (Elmegreen & Elmegreen 2006; Law et al. 2012; Conselice 2014). A minor-merger triggered face-on spiral at $z = 2.18$ remains the only thoroughly studied case in the literature (Law et al. 2012). Previous data show that spiral arms are less well-developed and more chaotic beyond $z = 0.5$ (Abraham & van den Bergh 2001). The onset of spiral structures in galaxies is proposed to occur at $z \sim 1.8$, when disks have developed a cool stellar component in a rotation-dominated disk (Elmegreen & Elmegreen 2014). However, indirect observational evidence shows that the Hubble sequence may already be in place at $z \sim 2.5$ (Wuyts et al. 2011), implying an even earlier onset of spiral arms.

Observational difficulties in identifying spiral arms at high redshift might have contributed to the rarity of spiral galaxies at $z \gtrsim 2$. In the local universe, spiral arms are visually classified through qualitative inspection of morphologies (e.g., Hubble 1926; Reynolds 1927; Sandage 2005). This visual classification scheme suffers strongly from observational biases caused by degraded resolution, cosmological surface brightness dimming, band shifting and imaging depth at high redshift (Abraham et al. 1996a; Giavalisco et al. 1996; Hibbard & Vacca 1997; Conselice et al. 2000, 2011). For example, spiral features can only be reliably identified for ~ 270 galaxies in the Hubble Ultra Deep Field (UDF) that have major axes larger than 10 pixels (Elmegreen et al. 2005). Most galaxies at $z \gtrsim 2$ have half-light radii of $\leq 0''.3$ (Allen et al. 2016), making the identification and quantification of spiral features such as the number of arms and pitch angles challenging (Shields et al. 2015). More studies like van den Bergh et al. (2002) focusing on testing the visibility of spiral arms at high redshift are required before concluding the actual number density of spiral galaxies at $z > 1$.

In addition to observational biases, cosmological simulations suggest a few physical processes at $z \gtrsim 2$ that hinder the formation of spiral arms. Earlier cosmological simulations suggest that the paucity of spiral galaxies at $z \gtrsim 2$ can be ascribed to high merger rates (Hammer et al. 2009), high gas accretion rates and multiple low angular momentum inflow cold streams (Cen 2014). Some cosmological simulations report that grand-design spirals are in place by $z \sim 3$ (Fiacconi et al. 2015) and the high-redshift spirals most likely originated from swing amplifications triggered by satellites. Detailed observations of spiral galaxies at high redshift will enable meaningful investigations into these physical processes that remain

elusive in simulations.

Gravitationally lensed galaxies with adaptive-optics (AO) aided observations have pioneered the measurement of physical properties of high-redshift galaxies (e.g., Swinbank et al. 2007; Stark et al. 2008; Jones et al. 2010a; Yuan et al. 2011, 2012, 2015; Swinbank et al. 2015). The lensing magnification allows selection of less massive systems and measurements on smaller physical scales than magnitude-limited studies. The spatial resolution in gravitational lensing observations can reach a few times 10-100 pc, important for minimizing beam-smearing effect and resolving star clusters (e.g., Jones et al. 2010b; Yuan et al. 2013a; Livermore et al. 2015; Vanzella et al. 2017; Rigby et al. 2017). In this paper, we report the integral-field spectroscopic (IFS) observation of a gravitationally lensed spiral galaxy *A1689B11* at $z = 2.54$. It is the second spectroscopically confirmed spiral galaxy at $z \gtrsim 2$ and is ~ 10 times less massive than the spiral galaxy of Law et al. (2012).

This paper is organized as follows. Section 2 describes our IFS observation, data analysis and lens models. Section 3 describes our results from the IFS data and derived physical properties of *A1689B11*. In Section 4 we discuss the nature of *A1689B11* and compare it to other galaxy samples. We summarize and conclude in Section 5. Throughout this paper, we adopt a standard Λ CDM cosmology with $\Omega_M=0.3$, $\Omega_\Lambda=0.7$ and $H_0=70 \text{ km s}^{-1} \text{ Mpc}^{-1}$. At the redshift of $z = 2.54$, $1''$ corresponds to a physical scale of 8.2 kpc.

2. OBSERVATIONS AND DATA REDUCTION

2.1. *Spiral Galaxy Candidate A1689B11*

The spiral galaxy candidate (hereafter *A1689B11*) was first recorded as “source 11” with a photometric redshift of $z = 2.9 \pm 0.2$ in the strong-lensing analysis of the galaxy cluster Abell 1689 (Broadhurst et al. 2005). Source 11 is gravitationally lensed into two highly magnified images B11.1 and B11.2, and a central image B11.3 (Figure 1). Broadhurst et al. (2005) was the first to point out the spiral feature of *A1689B11* and speculated it would be the highest redshift spiral galaxy if confirmed. The first spectroscopic redshift of $z = 2.5$ was provided in Limousin et al. (2007) based on rest-frame UV spectroscopy with Keck/LRIS. The detection of $H\alpha$ lines at $z = 2.54$ was reported in the near-infrared (NIR) multi-slit survey of Yuan et al. (2013b). This work presents the NIR IFS observation of image B11.1 (Figure 1). The imaging and photometric data used in this work are obtained from the Hubble Space Telescope (HST; proposal IDs: 11802, 9289, 11710) and the Spitzer Space Telescope archives (program ID: 20439).

2.2. *NIFS/Gemini Observations and Data Reduction*

AO aided NIR integral field spectroscopic observations with NIFS (Near-infrared Integral-Field Spectrograph; McGregor et al. 2003) were conducted between March 2013 and March 2014 on the Gemini North telescope under excellent weather conditions (average seeing $\sim 0''.5$, airmass 1-1.6). Our total allocated observational time was 13.5hrs in band 1 (program ID: GN-2013A-Q-23-64). The observation implemented a dithering pattern of “ABAABA” or “AB”, i.e., ~ 33 -50% of the target exposure time was spent on sampling sky frames in order to

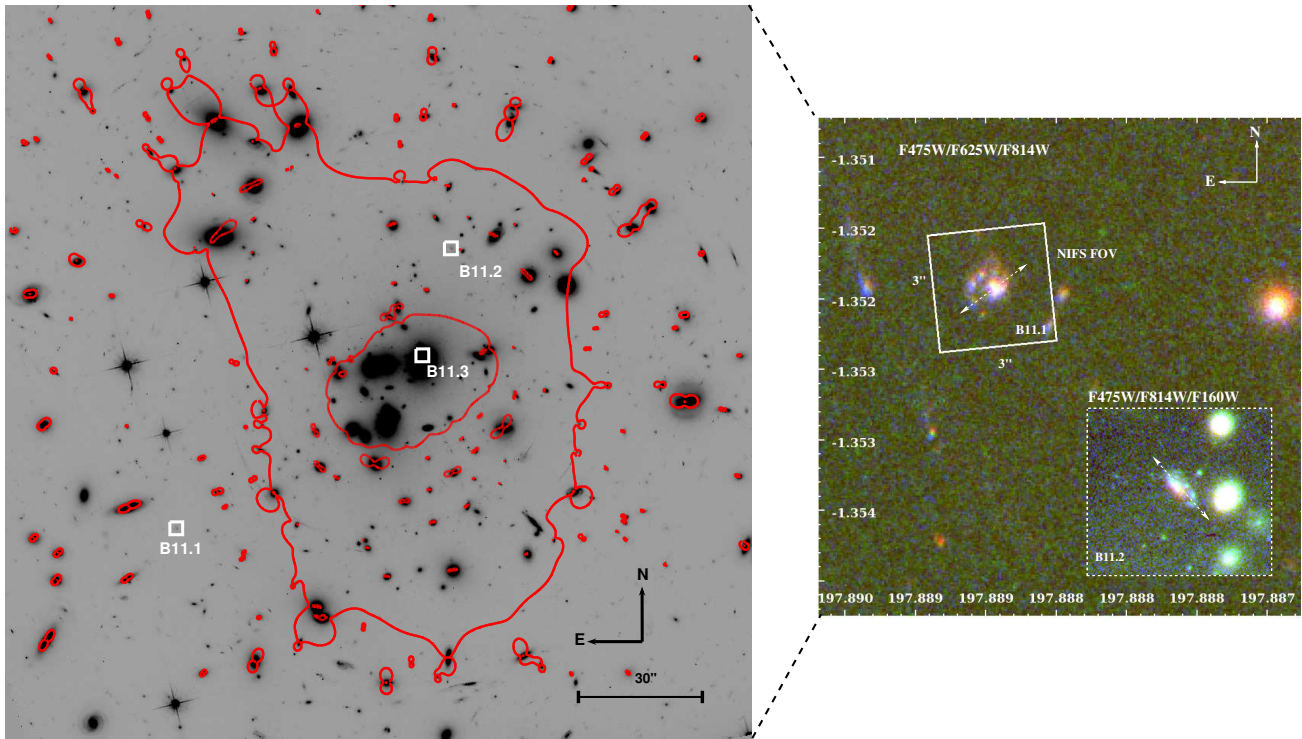


FIG. 1.— Left: The HST F814W image of the strong lensing cluster Abell 1689. Red lines show the critical lines at $A1689B11$'s redshift of $z = 2.54$ based on the lens model of Limousin et al. (2007). The outer red line is the tangential critical line; the inner red line is the radial critical line. White boxes show the positions of three multiple lensed images of $A1689B11$. B11.1 and B11.2 are well-resolved magnified images; B11.3 is the demagnified central image overlapping with one of the brightest cluster members. Right: The zoom-in HST 3-color (F475W/F625W/F814W) image (B11.1) of the lensed spiral galaxy $A1689B11$. The embedded 3-color (F475W/F814W/F160W) image (B11.2) shows extra color information from the HST WFC3 IR data. The dashed arrow denotes the major axis of the galaxy. B11.1 is stretched more in the minor-axis direction whereas B11.2 is stretched roughly equally in all directions by lensing. B11.1 and B11.2 show identical spiral morphology on both the image plane and the source plane in all available HST broad-band images (Appendix). Our NIFS/Gemini observation (solid white box) is centered on image B11.1 because of the laser guide star requirements.

facilitate a good sky background subtraction. The observation centered on image B11.1 instead of B11.2 because of the laser-guide star requirement of NIFS. The coordinates of the pointing center are given in Table 1. The field of view of NIFS is $3''.0 \times 3''.0$, with 29 slitlets, each $0''.1$ wide. NIFS delivers a spectral resolving power of $R \sim 5300$ in the K band, corresponding to a rest-frame Gaussian velocity resolution of $\sigma \sim 24 \text{ km s}^{-1}$.

The data were reduced using a varied Gemini IRAF package following standard procedures (e.g., Storchi-Bergmann et al. 2009). This IRAF package produces sky-subtracted, telluric-corrected and flux-calibrated datacubes. Each individual 900s exposure was spatially aligned based on the brightest $H\alpha$ spaxels and co-added based on a mean sigma-clipping procedure. The 1σ error datacubes were generated during the sigma-clipping process. We used the star HIP67004 of spectral type A0V for telluric correction and flux calibration. The systematic uncertainties in the flux calibration is estimated to be within 20%. A total of 27 ks on-source exposures were obtained in the K-band with an angular resolution of $\sim 0.1''$, corresponding to a median physical scale of ~ 400 pc on the source plane of $A1689B11$. Our observation reached a 3σ $H\alpha$ emission line surface brightness depth of $3 \times 10^{-17} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ arcsec}^{-2}$.

The astrometry of the datacube is calibrated by assigning the HST coordinate of the galaxy center to the bright-

est spaxel in the wavelength-collapsed 2-dimensional (2D) NIFS image and rotating with the positional angle from the observation. The astrometric uncertainty of our IFS datacube is about one spaxel, i.e., $\sim 0''.1$.

2.3. Emission Line Fitting

Our K-band observation was originally planned to detect $H\alpha$ and $[\text{N II}]$ lines. To analyze the IFS spectra, we first collapse the datacube in the wavelength dimension in the vicinity of the $H\alpha$ line to create an $H\alpha$ 2D map. We use the raw $H\alpha$ 2D map to generate an initial 2D mask that flags spaxels with no obvious $H\alpha$ line detections. We then manually inspect the spectrum of each individual spaxel and refine the mask with three types of visual flagging: (1) significant emission lines, (2) possible emission lines, and (3) no obvious emission lines. The visual inspection of the datacube is necessary in order to reject spaxels that are spurious for low signal-to-noise ratio (SNR) data. We use this visual mask to supervise our subsequent automatic emission line fitting procedure, i.e., we demand the fitting result to be consistent with our visual mask. For example, an emission line fitting result of $\text{SNR} \geq 5$ should have a visual flag of (1), whereas a line detection of $\text{SNR} < 3$ should be consistent with a flag type of (3). Those with $3 \leq \text{SNR} < 5$ should match the flag type of (2).

Our automatic emission line fitting procedure involves

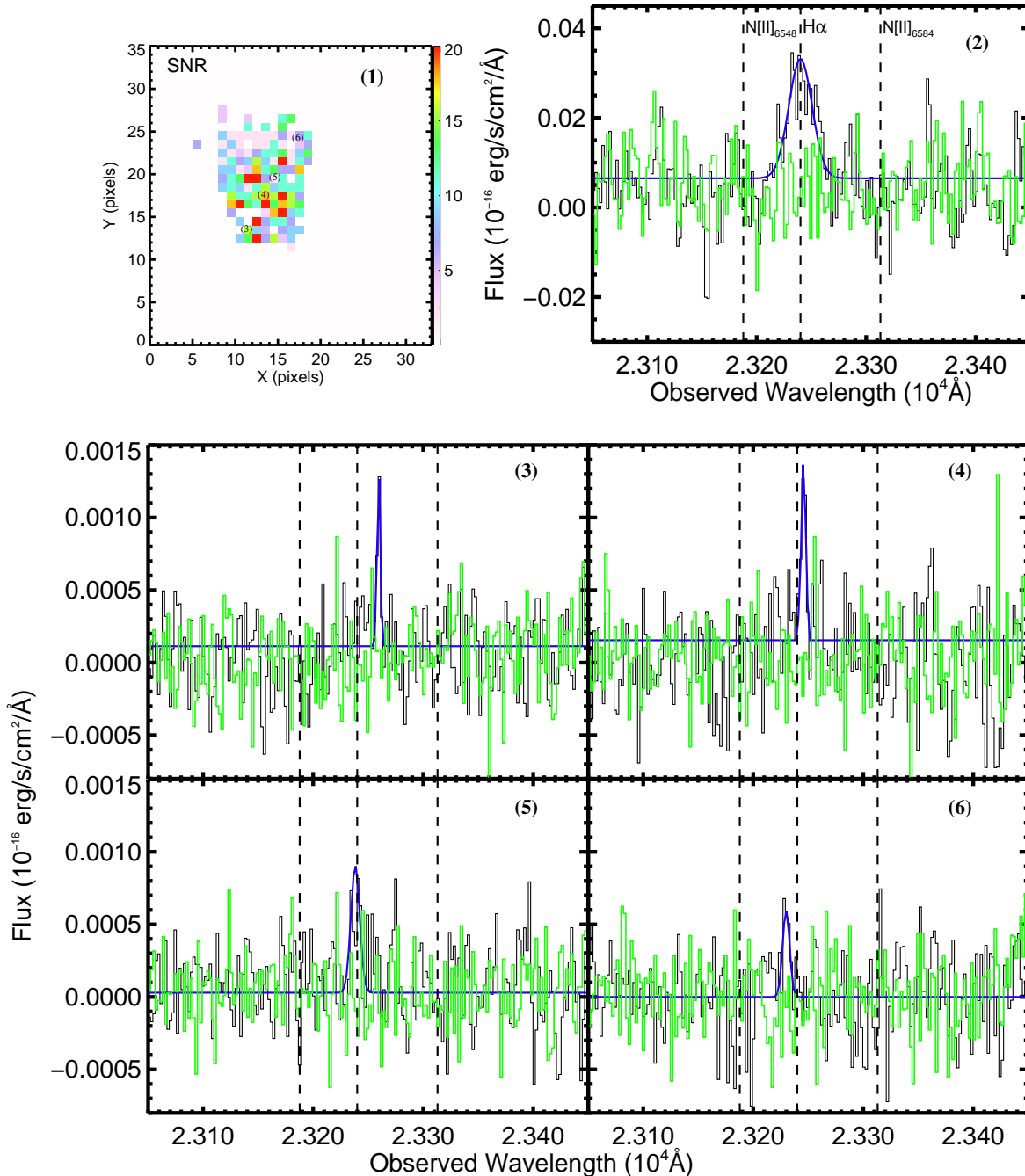


FIG. 2.— Examples of NIFS 1D spectra and 2D SNR map in the observed frame. Panel (1): 2D SNR map from H α emission line fitting (1 pixel corresponds to $0''.1$). Panel (2): Integrated 1D spectrum from coadding all NIFS spaxels, weighted by SNR. Panel (3)-(6): 1D spectrum from individual spaxels, chosen along the major axis of the galaxy to represent a range of SNR and velocity centroids. The locations of the corresponding spaxels are mark on panel (1). For panels (2)-(6), raw spectra are shown as black lines; the RMS of the sky residuals are shown in green; the best-fit Gaussian profiles for H α lines are shown in blue; vertical dashes lines indicate the expected positions of the [N II] λ_{6548} , H α λ_{6583} , and [N II] λ_{6583} emission lines at the kinematic center of the galaxy. [N II] lines are neither detected in any spaxel nor the integrated 1D spectrum. This figure highlights that H α emission lines from individual spaxels are genuinely narrow, even without beam-smearing corrections. The widths of the emission lines are approaching the instrumental resolution of NIFS (~ 2.5 wavelength channels here).

fitting Gaussian profiles simultaneously to three emission lines in each spaxel: [N II] λ_{6548} , 6583 and H α (Figure 2). The line profile fitting is conducted using a χ^2 minimization procedure weighted by the inverse of the variance spectrum. The fitting result is very sensitive to the weighting used. The 1σ error datacube generated from

the mean sigma-clipping procedure described in Section 2.2 does not provide good weighting to the emission line fitting, as it results in the SNR of the emission line being over-estimated and inconsistent with our visual mask flags. Instead we use the variance over a sky region devoid of emission lines as the weighting. We select a range

of sky spaxels to generate a variance spectrum (green lines in Figure 2) for each spaxel using a bootstrap procedure. The SNR of the emission line fitting result is consistent with the visual mask. This weighting method is similar to that used in Leethochawalit et al. (2016) and is reasonable because the variance of the NIR data is usually dominated by sky residuals.

H α lines are detected in ~ 130 individual spaxels at $\geq 5\sigma$ level (Figure 2, (1)). [N II] lines are not detected above 3σ in either individual spaxels or the integrated spectrum (Figure 2, (2)). Yuan et al. (2013b) report an [N II] line detection on one of the four slits configured with different positional angles on B11.1 and B11.2, implying a spatial variation of metallicities. We extract a mock slit spectrum from our NIFS data based on the slit setup of Yuan et al. (2013b), but no [N II] line is detected above 3σ . Deeper IFS data from our ongoing OSIRIS/Keck observation on image B11.2 will help to discern [N II] lines and the spatial metallicity distribution of *A1689B11*. Figure 2 shows the H α 2D SNR map and examples of single Gaussian fits to H α emission lines of individual spaxels. An integrated spectrum from coadded spaxels is also presented in Figure 2 for comparison. Figure 2 highlights the necessity of high-spatial resolution in order to distinguish beam-smearing from intrinsic line width. It is also apparent from Figure 2 that the H α lines detected on individual spaxels have very narrow widths (approaching the NIFS instrumental line width) and show systematic offsets in the centroids. The focus of this paper is to report the kinematics of H α lines (see Section 3).

2.4. Gravitational Lens Modeling

Abell 1689 is one of the most studied lensing clusters with well-constrained mass models. We use the best-fit model from Limousin et al. (2007) and the software *Lenstool*² (Kneib et al. 1993; Jullo et al. 2007) to reconstruct the source-plane properties of *A1689B11*.

We show the lensing configuration and critical lines of *A1689B11* in Figure 1 (left). *A1689B11* is lensed into a three-image system: B11.1, B11.2 and B11.3. Image B11.2 lies in between the tangential and the radial critical line, resulting in a large flux magnification factor ($\mu = 12 \pm 2$) and a relatively undistorted image (i.e., equally magnified in the major and the minor axis). Our NIFS observation is centered on the less magnified ($\mu = 7.2 \pm 0.8$) image B11.1. B11.1 is magnified more in the minor axis direction, making it appear more face-on than the intrinsic image. The geometric magnification for image B11.1 is ~ 4 (along the minor axis) $\times 1.8$ (along the major axis). A demagnified central image B11.3 is predicted by the lens model and also observed to be overlapping with one of the brightest cluster members of Abell 1689 (Broadhurst et al. 2005; Limousin et al. 2007).

The statistical error of the lens modeling is estimated from a set of Markov Chain Monte Carlo (MCMC) realizations implemented in *Lenstool* and is typically $\sim 10\%$ for Abell 1689. Both multiple images (B11.1 and B11.2) yield consistent source-plane morphologies and positions within $0.2''$. Because B11.1 and B11.2 are not close to the critical lines, the systematic errors in the source-plane

morphologies are small compared to the case of giant arcs. Our source-plane reconstructed morphologies are robust within the model statistical uncertainties. The NIFS datacube is reprojected to the source plane after being remapped to the HST coordinates using the calibrated astrometry. Emission line fitting analyses have been carried out on both the image and source plane and the results are self-consistent. We present both the image-plane and source-plane properties in the following sections.

3. ANALYSIS AND RESULTS

3.1. NIFS Data Analysis

3.1.1. NIFS H α Map and Kinematics

In Figure 3 we show the 2D H α flux intensity map and 2D H α kinematic maps on the observed frame (image plane), i.e., before lensing reconstruction. To facilitate the comparison with HST images, we mark the extent of HST broad-band detections in contours in all panels. We use the HST/F814W band image as an example because it has the deepest exposure among all HST filter observations. We include the morphology of other HST bands in Section 3.2.

In order to preserve the spatial resolution offered by lensing and because we have sufficient detections on individual spaxels, we do not attempt to bin the original IFS datacube. The NIFS maps in Figure 3 are un-smoothed and include only data with $\text{SNR} \geq 5$. H α emission lines are detected throughout the HST contours; H α emission lines are relatively stronger near the galaxy core and star forming clumps in the spiral arm.

The rest-frame line-of-sight velocity 2D map is derived from the H α line centroids with respect to the systematic redshift and the velocity dispersion map from the H α Gaussian line width. We use the median H α line center to calculate the systematic redshift. The instrumental profile has been subtracted in quadrature from the best-fit Gaussian width to derive the intrinsic line width. Figure 3 shows a systematic rotation and uniformly small velocity dispersion across the disk on the observed plane.

The intrinsic morphology and NIFS measurements are shown in source-plane maps in Figure 4. We use a grid subsampling of *echant* = 20 on the image-plane and *s_echant* = 10 on the source-plane in *Lenstool* to optimize the spatial resolution of the source-plane reconstruction. The source-plane HST ACS image in Figure 4 is unbinned and has a pixel scale of $0''.005$ after subsampling. The source-plane NIFS datacube is rebinned adaptively by 5-11 pixels to allow for H α line fitting of $\text{SNR} > 5$ in each bin. Note that because of the difference in the point spread function (PSF) on the image and source planes, the improvement in the source-plane SNR is usually less than what the magnification map predicts. On average we achieve a source-plane spatial resolution of ~ 150 pc on the HST ACS images and ~ 400 pc on the NIFS 2D maps.

The source-plane 2D velocity map in Figure 4 clearly shows a velocity gradient consistent with a systematic rotation. The velocity dispersion is enhanced slightly in the kinematic center due to beam-smearing from the rotation. We estimate a maximum beam-smearing effect at the kinematic center to be $\sim 24 \text{ km s}^{-1}$ based on the

² <https://projets.lam.fr/projects/lenstool/wiki>

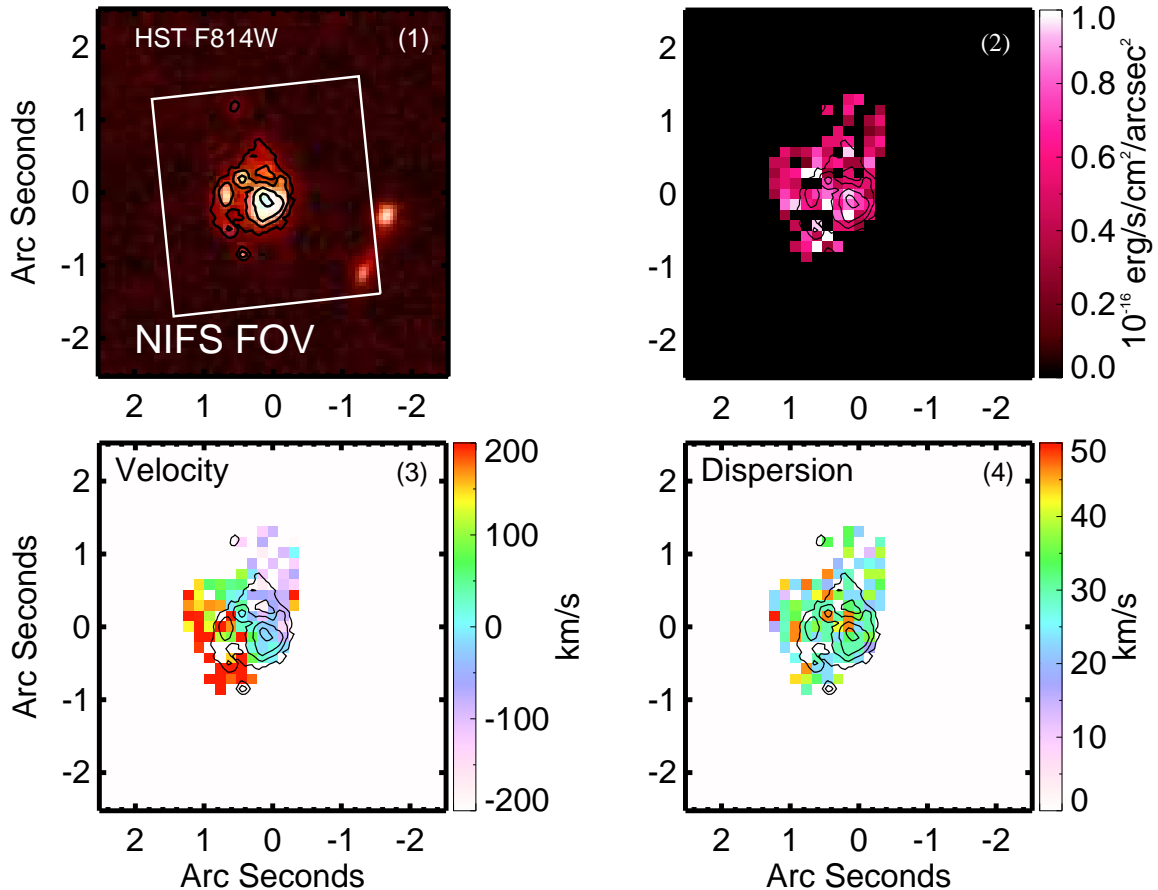


FIG. 3.— HST morphology and NIFS 2D maps on the image-plane. Panel (1): The HST F814W morphology of observed image B11.1. Black contours outline different surface brightness levels of the F814W image, including star forming clumps in the core and on the spiral arms. The white box shows the NIFS FOV ($3''$ by $3''$). (2): NIFS H α intensity 2D map. (3): NIFS H α velocity 2D map. (4): NIFS H α velocity dispersion 2D map. Black contours in (2)-(4) are the same as panel (1) and have been astrometrically aligned with the NIFS observation. All NIFS 2D maps are presented with the observed spaxel scale without smoothing/binning. Only data with $\text{SNR} \geq 5$ are included.

velocity map of panel (3) in Figure 4, consistent with the 2D velocity dispersion map of panel (4). We do not see any significant spatial correlation of the H α intensity map with the velocity dispersion map. Figure 4 confirms that *A1689B11* has a uniformly low velocity dispersion across the disk on the source-plane. The mean velocity dispersion averaged over all spaxels is $V_{\sigma, \text{mean}} = 23 \pm 4 \text{ km s}^{-1}$. Excluding the central spaxels that are affected most by beam-smearing, the average velocity dispersion on the outer disks is $V_{\sigma, \text{outer-disk}} = 15 \pm 2 \text{ km s}^{-1}$.

3.1.2. Disk Model Fitting

We use an empirically motivated arctangent function from Courteau (1997) to model the 2D velocity field:

$$V(R) = V_0 + \frac{2}{\pi} V_c \arctan \frac{R - R_0}{R_t}; \quad (1)$$

The source-plane line-of-sight velocity $v_s(R)$ is related to the intrinsic velocity $V(R)$ by the inclination angle i :

$$v_s(R) = V(R) \sin(i); \quad (2)$$

The source-plane radius vector R_s is related to the in-

trinsic radius vector R by:

$$\vec{R}_s = \vec{R} \cos(i) \begin{bmatrix} \cos(PA), & -\sin(PA) \\ \sin(PA), & \cos(PA) \end{bmatrix}, \quad (3)$$

where PA is the positional angle. For the definition of PA and i , we use the same convention as GALFIT (Peng et al. 2010), i.e., $PA = 0$ is to the north (up), and $PA = 90$ to the east (left); $i = 0$ is face-on and $i = 90$ is edge-on.

The seven free parameters of the 2D disk model are thus: central velocity V_0 , inclination i , position angle PA , disk dynamic center \vec{R}_0 (R_{0x} , R_{0y}), turn-over radius R_t , and asymptotic velocity V_c . Note that V_0 is close to zero if the median H α line center is a good approximation for the central velocity. The best-fit model is obtained using a χ^2 minimization procedure similar to that described in Jones et al. (2010b). We use the statistical 1σ errors from the emission line fitting to compute the χ^2 . The uncertainties of the best-fit parameters are estimated by perturbing the model until the χ^2 increases by one standard deviation from the best-fit model.

Figure 5 shows the best-fit 2D disk model and the residual maps. The reduced χ^2 of our best-fit model is 2.1 and the best-fit parameters are: asymptotic velocity $V_c = 200$

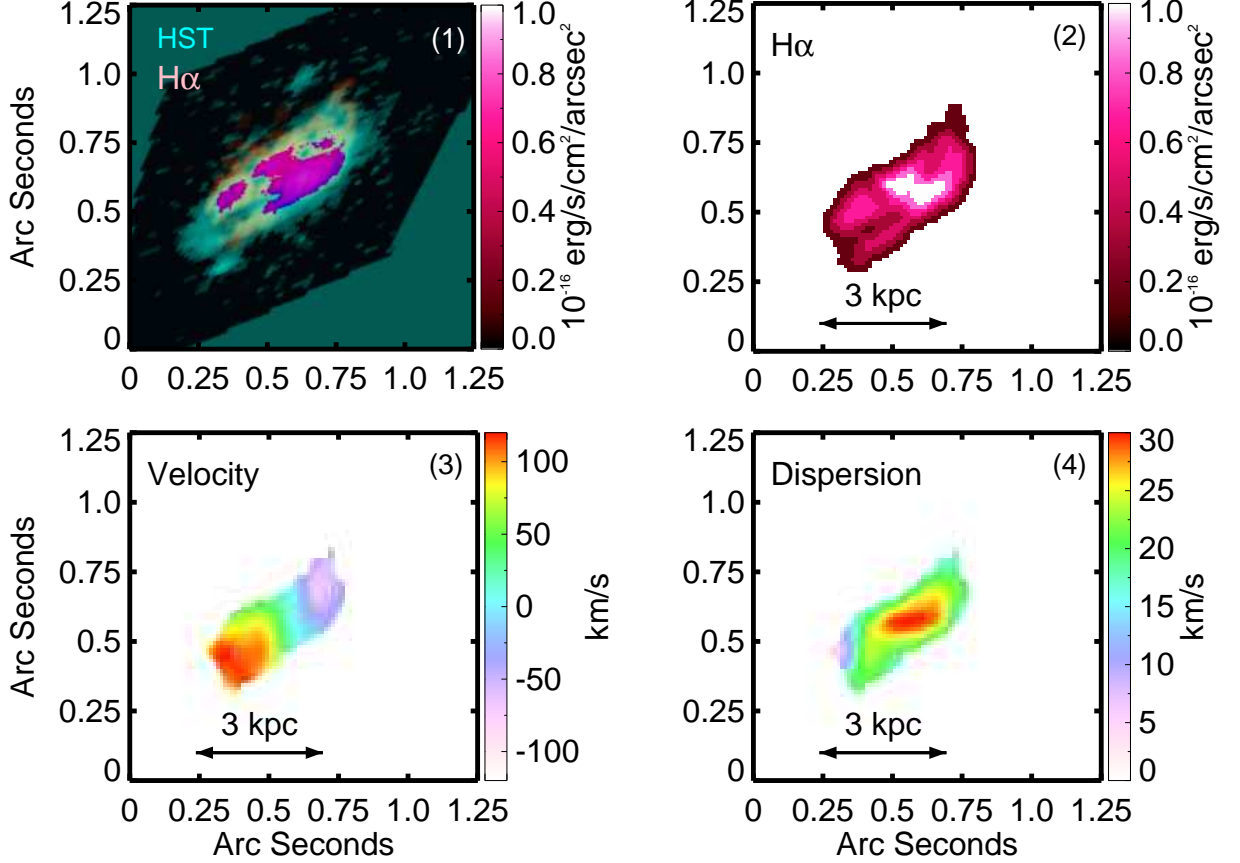


FIG. 4.— The HST morphology and NIFS 2D maps on the source-plane. Panel (1): The NIFS $H\alpha$ image (pink) on top of the HST F814W source-plane image (cyan). For the HST image, we used a subsampling of 20 on the image-plane and 10 on the source-plane to optimize the spatial resolution of the source-plane reconstruction. The pixel scale of the source-plane HST image is therefore $0''.005$. The source-plane NIFS datacube is re-binned adaptively by 5-10 to allow for $H\alpha$ SNR > 5 in each bin. (2): NIFS $H\alpha$ intensity 2D map. Note that the pixel scale of the source-plane NIFS image is $0''.01$. The $H\alpha$ image in panel (1) is astrometrically aligned and re-binned to match the HST F814W source-plane resolution. (3): NIFS $H\alpha$ velocity 2D map. (4): NIFS $H\alpha$ velocity dispersion 2D map.

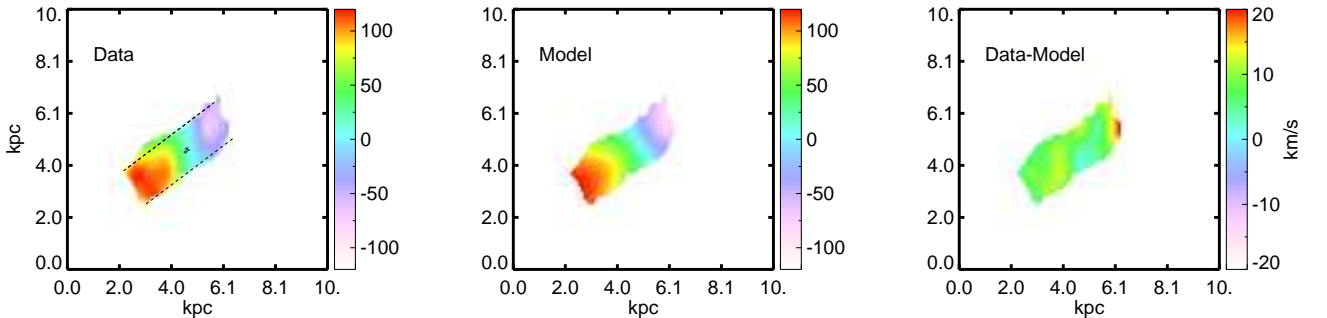


FIG. 5.— 2D disk model fit to the velocity field. The left panel is the data. The best-fit kinematic center is marked as a cross. The kinematic center matches the center defined by HST broad-band photometry (shown as the star symbol) within $0''.1$. The dashed line displays a slit used to extract the 1D velocity in Figure 6. Middle panel: best-fit model. Right panel: residual.

$\pm 12 \text{ km s}^{-1}$, turn-over radius $R_t = 1.7 \pm 0.1 \text{ kpc}$, inclination angle $i = 51 \pm 2$ degrees, and positional angle $PA = -37 \pm 2$ degrees. The kinematic center is marked on Figure 5 and is consistent with the brightest intensity from both the $H\alpha$ image and the HST broad-band images within $0''.1$.

We can also use Eq. (1) to fit the 1D rotation curve.

We extract the 1D rotation velocities using a $0''.8$ slit along the major axis with the PA and the kinematic center from the best-fit 2D model (see Figure 5, left panel). The seven-parameter 2D model then reduces to a five-parameter 1D model. In Figure 6, each data point represents one resolution element ($0''.1$) binned along the slit and the error bars are the standard deviation of each

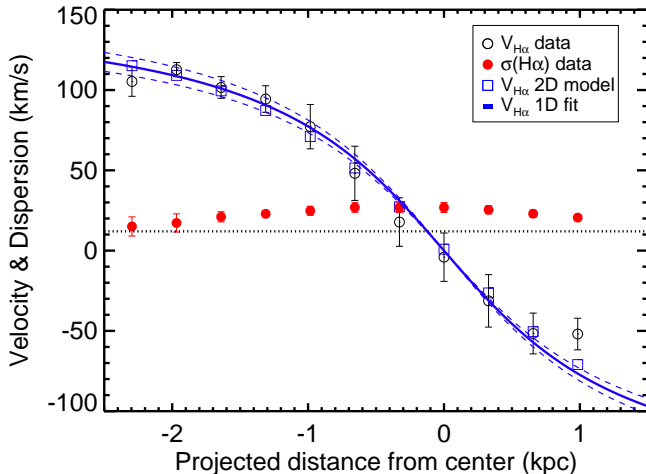


FIG. 6.— One-dimensional $H\alpha$ velocity (black empty circles) and velocity dispersion (red filled circles) measured on a slit overlaid along the major axis (Figure 5). Each data point represents one resolution element along the major axis. The error bars are standard deviations of spaxels binned perpendicular to the major axis. The blue boxes are data extracted from the 2D best-fit model. The best-fit 1D rotation curve and its 1σ variation are shown as blue solid and dashed lines. The horizontal black dotted line highlights the thermal broadening threshold of the $H\alpha$ line.

data bin. We show the best-fit 1D rotation curve and its 1σ variations in Figure 6. We find that the best-fit parameters from the 1D model are consistent with the 2D disk model within statistical errors. Note that we have also tested fitting a 3D disk model such as Di Teodoro & Fraternali (2015) to the datacube. Unfortunately, we do not have sufficient SNR per spaxel for a reliable fit of the 3D model. Our 2D and 1D disk model provides a reasonably good fit to the data. Overall, a regular rotating disk model is sufficient to explain the velocity field of *A1689B11*. The small residuals on the edge of the disk (Figure 5 and Figure 6) are most likely due to under-estimation of the observational noise.

The simple 2D disk model does not consider beam-smearing and variations of the velocity dispersion. However, as discussed in Section 3.1.1, the disk shows a uniformly low velocity dispersion and the beam-smearing does not contribute significantly to the shape of the velocity field. We show the 1D velocity dispersions (red points) along the kinematic major axis in Figure 6. The velocity dispersion in the galactic center is mildly enhanced due to beam-smearing effects. The dispersion on the outer disk is approaching the thermal broadening threshold (12 km s^{-1}) of $H\alpha$ emission line (black dotted line).

We discuss the kinematics of *A1689B11* and compare it with other samples in Section 4.

3.2. HST morphologies

To compare our NIFS data with HST images in more detail, we analyze source-plane morphologies for all HST ACS and WFC3 bands. B11.1 and B11.2 are well-detected in all four IRAC bands from Spitzer. We do not gain resolved morphological knowledge from IRAC because of its poor spatial resolution ($\sim 2''.5$, larger than the lensed images). Therefore IRAC data is only used when deriving global properties of *A1689B11* in Section 3.3.

We take full advantages of the multiple images from

gravitational lensing and combine information from both images B11.1 and B11.2 to derive the best source-plane morphology. B11.1 is covered by 6 filter bands (F475W, F625W, F775W, F814W, F105W, F140W) and B11.2 is covered by these bands plus the additional 2 bands of F125W and F160W. Source-plane morphologies for individual HST bands are reconstructed using the same method as implemented for the F814W image described in Section 3.1. The spiral features are more prominent in the HST/ACS optical images ($0''.05$ resolution), whereas HST/WFC3 images ($0''.1$ resolution) show mostly the central disk. We include source-plane morphologies for individual bands in the Appendix.

We perform GALFIT (Peng et al. 2010) on the source-plane images using a single exponential disk model. The fitting yields consistent morphological parameters for all bands of HST images and for both B11.1 and B11.2. The mean and standard deviation of the scale lengths (r_s) calculated from all 14 source-plane images are $r_s = 1.3 \pm 0.4$ kpc, with an inclination angle of $i = 55 \pm 10$ degrees and $PA = -36 \pm 6$ degrees. The single exponential disk model from GALFIT are in excellent agreement with geometric parameters derived from the NIFS kinematic 2D disk model (summarized in Table 1). For the convenience of comparing with various definitions of radius in literature, we convert r_s to the half-light radius $R_{1/2}$ and the effective radius R_e using the empirical relation of $R_{1/2} = 2.2r_s$, $R_e = 1.68r_s$ (Glazebrook 2013). The half-light radius of *A1689B11* is therefore 2.6 ± 0.7 kpc. The central area of *A1689B11* shows significant $H\alpha$ emission from our NIFS data and some elongated substructures in the HST ACS images (Figure 4; Appendix). The central area could be a superposition of a star cluster and a bulge/bar component. We experimented with adding a bulge component to the exponential disk, however the fitting does not converge for most of the images. We are currently investigating a more sophisticated procedure for the bulge-disk decomposition and our preliminary result shows a very small bulge component. We will report the full analysis in a future work focusing on the pitch-angle and bulge correlation of this spiral galaxy (Yuan, in preparation).

Because gravitationally lensed images cover an extended area, the probability of having contaminated foreground/background sources in the field of view of lensed images is larger than a non-lensed high- z galaxy case. We cross-compare substructures on the source-plane images of B11.1 and B11.2 to reject clumps/knots that may not be associated with the spiral galaxy. We use the F814W HST images that have the deepest exposure and best spatial resolution for the cross-examination. Figure 7 demonstrates our clump identification and rejection procedure. We first manually identify bright clumps on the right panels of Figure 7. There are 7 clumps (c0-c4, x1, x2) marked for lensed image B11.1 and 9 clumps (c0-c4, x3, x4, x5, x6) for lensed image B11.2. We then predict the source-plane positions and brightnesses of the clumps (c0'-c4', x1'-x6') and cross-examine them in the left panels of Figure 7. Because lensing conserves surface brightness and images B11.1 and B11.2 cover similar intrinsic areas of the lensed galaxy, the resolved source-plane substructures of B11.1 and B11.2 should be iden-

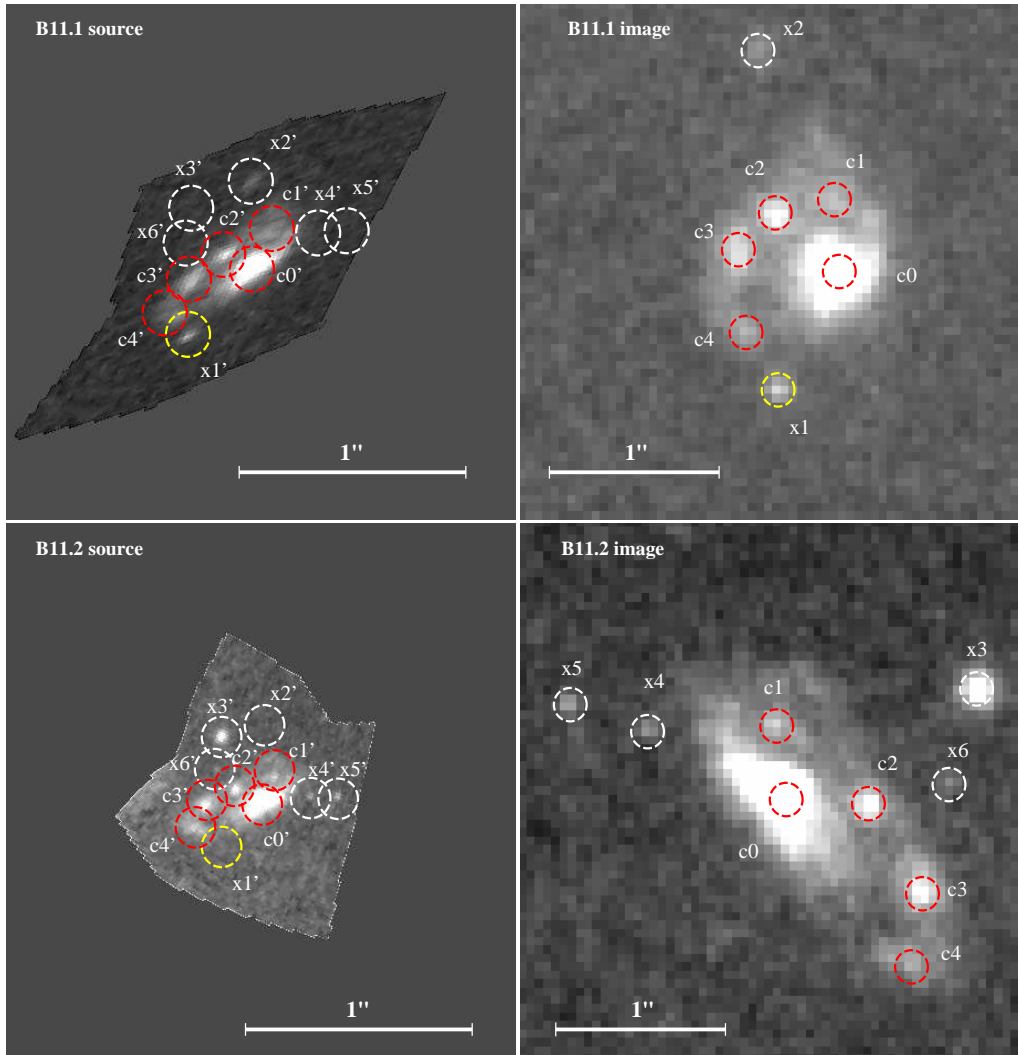


FIG. 7.— Rejecting substructures that are not associated with the spiral disk by cross-comparing substructures on the image and source plane of B11.1 and B11.2. Top left panel: source-plane morphology of B11.1. Top right: image-plane morphology of B11.1. Bottom left: source-plane morphology of B11.2. Bottom right: image-plane morphology of B11.2. All images are based on the HST/F814W band. The circles on the right panels are manually identified clumps on the image-plane. Corresponding positions of these clumps on the source-plane are predicted by lens models. Clumps that are detected ($\text{SNR} > 5$) in both of the left panels are bona fide substructures (red circles) of the spiral galaxy, and vice-versa (white circles). The yellow circle shows a clump that has a marginal ($\text{SNR} \sim 3$) detection in B11.2 but is rejected because of the inconsistency in brightness. All circles have a diameter of $0''.2$, representing the RMS of our lens model reconstructions.

tical from the same HST observation³. Source-plane clump brightness and positions that are not consistent within the uncertainties of lens models are considered as interlopers. For example, clump $x2'$ is identified initially on B11.1 but is not detected on B11.2 within the lensing position uncertainty ($\text{RMS}=0''.2$), we therefore reject clump $x2'$ and consider it as a foreground or background source. Similarly, clumps $x4'$ and $x5'$ identified initially on B11.2 are not detected on B11.1 and are rejected. For clump $x1'$ (yellow circles in Figure 7) that is initially identified on B11.1, there is a marginal 3σ detection in the source-plane of B11.2, however, the lens model predicts a source-plane brightness that is ~ 3 times brighter than what is observed. We therefore reject $x1'$ because

³ This method is not suitable in cases where the lensed images are crossed by critical lines, in which case the multiple lensed images may represent different parts of the intrinsic galaxy.

of the inconsistency in flux magnification uncertainties ($\sim 10\text{-}20\%$). Note that clumps/knots that are identified as foreground/background objects in this lensing analysis are at significantly different cosmological distances from the source galaxy and are therefore not satellites. The final combined morphologies of B11.1 and B11.2 are shown in the Appendix.

Finally, we use colors of the clumps as an alternative check for substructure identities. We measure aperture photometry for all clumps in B11.1 in Figure 7 using HST broadband images. A reliable ($\text{SNR}>5$) photometry can only be measured for six clumps across a minimal of 3 bands: $c0'$ - $c4'$ and $x2'$. We then carry out Spectral Energy Distributions (SEDs) fitting for the clumps using the software LE PHARE (Ilbert et al. 2010). We fix the redshift at $z = 2.54$ and use the stellar population synthesis models of Bruzual & Charlot (2003). We choose an initial mass function (IMF) of Chabrier (2003) and

the Calzetti et al. (2000) attenuation law, with $E(B - V)$ ranging from 0 to 2 and an exponentially decreasing star formation history. The best-fit SEDs are shown in Figure 8. Clumps $c0'$ - $c4'$ show similar SEDs, whereas $x2'$ has a very different SED, consistent with our lensing source-plane position analysis that $x2'$ is most likely an interloper. We also show in Figure 8 the best-fit SED for the total photometry of B11.1 (Section 3.3). In total, we confirm 5 bona fide clumps ($c0'$ - $c4'$) for *A1689B11*. There are significant $H\alpha$ emissions detected on the central clump $c0'$ and on clump $c2'$ on the spiral arm; $H\alpha$ lines are also detected on clump $c1'$ and $c3'$ on the spiral arm (see Figure 4). A detailed analysis on individual clump properties will be reported in a separate work when we combine our ongoing OSIRIS/Keck data analysis on image B11.2. Our preliminary result shows that the sizes (400–600 pc) and surface star forming densities ($0.2 - 0.3 M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$) of the clumps are comparable to high-redshift SF galaxies and are in the intermediate range of what has been reported for $z > 1$ clumpy galaxies (e.g., Genzel et al. 2011; Jones et al. 2010b; Livermore et al. 2015).

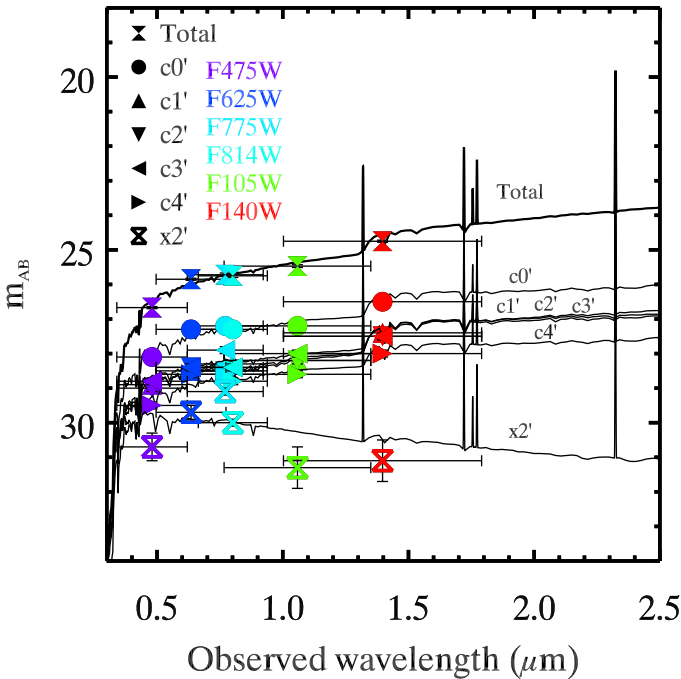


FIG. 8.— SED fitting to the source-plane photometry for the entire galaxy (total) and individual clumps ($c0'$ - $c4'$, $x2'$) identified in Figure 7. We show the six clumps that have reliable aperture photometry measured on the source-plane. Details of the SED fitting are described in Section 3.2. Clump $x2'$ has a very different SED shape compared to the total galaxy and the rest of the clumps, consistent with our lensing position analysis that $x2'$ is likely a foreground/background object.

3.3. Star formation Rate, Stellar and Dynamic Mass

We estimate the total stellar mass from the SED fitting of broadband photometries from the HST ACS and the Spitzer IRAC data (details in Yuan et al. 2013b) (best-fit SED shown in Figure 8). The best-fit stellar mass

for *A1689B11* is $M_{\text{star}} = 10^{9.8 \pm 0.3} M_{\odot}$ and the best-fit extinction value is $E(B - V)_{\text{stellar}} = 0.22$. The total dust-corrected SFR from the SED fitting is $\text{SFR}_{\text{SED}} = 22 \pm 3 M_{\odot} \text{ yr}^{-1}$. All values have been corrected for the lensing flux magnification. The dust-uncorrected SFR from the total $H\alpha$ fluxes of our NIFS observations is $\text{SFR}_{H\alpha, \text{nodust}} = 3.9 \pm 0.4 M_{\odot} \text{ yr}^{-1}$. Using the nebular dust extinction $E(B - V)_{\text{nebular}} = 0.73$ from (Yuan et al. 2013b) and the nebular attenuation curve of Cardelli et al. (1989), the dust-corrected SFR from $H\alpha$ is $\text{SFR}_{H\alpha} = 22 \pm 2 M_{\odot} \text{ yr}^{-1}$, in agreement with SFR_{SED} . The total stellar mass and SFR of *A1689B11* are consistent with a SF galaxy on the $z \sim 2$ mass-SFR relation (the main-sequence) within the ~ 0.3 dex 1σ scatter (e.g., Zahid et al. 2012) of the relation. The total SFR of *A1689B11* is 10-20 times higher than a typical spiral galaxy of similar masses at $z \sim 0$.

Because we do not have spatially resolved dust attenuation measurements, the following estimation of SFR surface density is indirect and based on a few assumptions. Assuming the spatially resolved $E(B - V)_{\text{nebular}}$ is similar to the global $E(B - V)_{\text{nebular}}$ measured from the slit data of Yuan et al. (2013b), then the average SFR surface density is $\Sigma_{\text{SFR}} = 0.3 M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$. This value of Σ_{SFR} is in the intermediate range of $z > 1$ SF galaxies and is 1-2 orders of magnitude higher than local SF galaxies (e.g., Swinbank et al. 2012; Fisher et al. 2017; Zhou et al. 2017). Assuming a simple Schmidt-Kennicutt (KS) relation (Kennicutt 1998), we convert Σ_{SFR} into a gas surface density of $\Sigma_{\text{gas}} \sim 158 M_{\odot} \text{ pc}^{-2}$. Assuming the surface area of the gas is $2\pi R^2$, where R is the radius where $H\alpha$ are detected (~ 1.7 kpc), we then derive a gas fraction ($f_{\text{gas}} = M_{\text{gas}} / (M_{\text{gas}} + M_{\text{star}})$) of $\sim 18\%$. The f_{gas} is in the lower range of star-forming galaxies at $z > 1$ but still significantly higher than local SF galaxies (e.g., Tacconi et al. 2013). We compute the Toomre Q -parameter (Toomre 1964) for a gas-dominated disk as defined by $Q \approx \kappa V_{\sigma} / \pi G \Sigma_{\text{gas}}$, where κ is the epicyclic frequency of the galaxy's rotation, V_{σ} is the gas velocity dispersion, and Σ_{gas} is the surface mass density. To compute κ , we assume a Keplerian disk so that $\kappa = \Omega$, where Ω is the angular frequency. The Ω calculated at the half-light radius is $\Omega \sim 2\pi / 100 (Myr)$. Adopting $V_{\sigma} = 23 \text{ km s}^{-1}$, we find $Q_{\text{gas}} \sim 0.7$ for *A1689B11*. The small Q_{gas} is commonly measured in clumpy high-redshift SF galaxies and is consistent with the scenario that large SF clumps form in the violent disk instability (Genzel et al. 2011; Law et al. 2012; Glazebrook 2013; Shibuya et al. 2016). We caution that the intrinsic Σ_{SFR} can easily change by a factor of two because of the unknown dust attenuation. The systematic errors related to the methodology of deriving f_{gas} and Q_{gas} are also highly uncertain and proper calculation requires estimation of the Mach number and molecular gas observations (e.g., Federrath et al. 2017b).

The dynamical mass assuming a rotationally supported disk is $M_{\text{dyn}} = R V_{\text{rot}}^2 / G = 10^{10.2 \pm 0.1} M_{\odot}$, where V_{rot} is the asymptotic velocity V_c and the radius R is the turnover radius R_t from the 2D disk model (Section 3.1.2). The ratio of the dynamical mass to the stellar mass for the inner ~ 2 kpc is therefore $M_{\text{dyn}} / M_{\text{star}} \sim 2.8$. Taken the M_{dyn} as the sum of dark matter and baryonic matter, then the inferred dark matter mass fraction within the inner ~ 2 kpc (compatible to $R_{1/2}$) is $f_{DM} = 60\%$.

The value of f_{DM} is in the typical range of local late-type spiral galaxies of similar V_c and much larger than the massive baryon-dominated clumpy disk galaxies at $z \sim 2$ (Genzel et al. 2017; Lang et al. 2017). However, this conclusion is subject to uncertain observational and methodological errors of f_{DM} . For example, by simply propagating the errors of M_{dyn} and M_{star} , we obtain $f_{DM} = 60\% \pm 40\%$. In addition, in order to decompose the contributions of the baryonic disk and the dark matter halo to the total rotation curve, assumptions on the mass-to-light ratio (M/L) and scale-height of the stellar disk have to be made (e.g., Aniyani et al. 2016), both are difficult to constrain for our galaxy. We therefore caution again that values derived in these two paragraphs are dominated by systematic errors and should be interpreted with some caution.

4. DISCUSSION

4.1. *Spiral arms vs merger*

The morphology of *A1689B11* is indicative of a prototype spiral arm but not conclusive. The clumpy morphology is also suggestive of a merger. However, by combining the morphology with the high-spatial resolution kinematics presented in this work, we rule out the merger scenario for *A1689B11*.

Distinguishing mergers from isolated disks is extremely tricky at high redshift. Neither morphological nor kinematic classifications alone can unambiguously exclude mergers. Morphological classifications of mergers rely on footprints of interactions such as bridges, tidal tails and double nuclei. These features can be easily missed at high redshift because of surface brightness dimming, size evolution and band shifting (e.g., Hibbard & Vacca 1997; Hung et al. 2015). High-redshift galaxies are clumpier and show more irregular structures than local galaxies, further complicating the visual characterization of mergers (Elmegreen et al. 2007). Kinematic classifications assume that isolated disks exhibit smooth velocity gradients whereas mergers show more asymmetric and chaotic kinematic features (Shapiro et al. 2008; Dicaire et al. 2008; Colina et al. 2005). However, this assumption does not consider post-coalescence mergers which may also display disk-like kinematics (Bellocchi et al. 2012). A ‘‘morpho-kinematic’’ classification that combines the morphological and kinematic criteria is proposed as a more robust approach (Rodrigues et al. 2017).

The velocity field of *A1689B11* is consistent with a rotating isolated disk based on current morphological and kinematic classification schemes. We first use the kinematic criteria of the SINS survey (Shapiro et al. 2008) and derive V_{asym} and σ_{asym} from the 2D velocity and velocity dispersion map. We find $V_{\text{asym}} = 0.1$ and $\sigma_{\text{asym}} = -0.08$, placing *A1689B11* in the isolated disk region of the V_{asym} versus σ_{asym} diagram. The unique clump identification method presented in Section 3.2 rejects most of the minor-merger like features as either foreground or background sources. For clumps that do associate with *A1689B11*, their velocities are consistent with rotating along with the main disk and their velocity dispersions do not show deviations from the rest of the disk, inconsistent with the scenario that the clumps are mergers. To further rule out the scenario of post-coalescence mergers, we apply the five ‘‘morpho-kinematic’’ criteria

of Rodrigues et al. (2017). We find that *A1689B11* satisfies all five criteria of an isolated rotating disc: (1) the velocity map has a single velocity gradient (Section 3.1.2 and Figure 5); (2) $V_{\text{rot}}/V_{\sigma} > 1$ ($= 9 - 13$ for *A1689B11*); (3) there is a V_{σ} -peak coinciding with the centre of rotation (Section 3.1.2 and Figure 4); (4) there is no mismatch between the kinematic and morphological PAs, i.e., $\Delta PA < 30^\circ$ ($\Delta PA = 1^\circ$ for *A1689B11*, Section 3.1.1 and 3.2; Table 1); and (5) The rotation center matches the stellar mass center within $0''.4$ ($\lesssim 0''.1$ for *A1689B11*; Section 3.1.1 and 3.2; Figure 5). Based on these arguments, we therefore exclude mergers as the origin of *A1689B11*.

4.2. *Spiral arms vs clumpy disks and irregulars*

Spiral galaxies and irregular galaxies are two distinct morphological classes in the local universe (e.g., Hubble 1926). However, as the general morphology of distant galaxies becomes more chaotic and irregular (e.g., Abraham et al. 1996a,b; Conselice et al. 2005; Elmegreen et al. 2007; Shibuya et al. 2016), spirals and clumpy/irregular galaxies do not have to be mutually exclusive at high redshift. For example, studies focused on spiral morphologies show that spiral structures are highly disturbed and arms are less well-developed at $z \gtrsim 0.5$ (Abraham & van den Bergh 2001; Elmegreen et al. 2005). The less well-defined spiral structures at high redshift could be caused by either intrinsic evolutions and/or observational biases.

Spiral arms are the main sites of star formation in the local universe. The brightness, sizes and surface densities of star forming regions are much larger at high redshift (e.g., Elmegreen et al. 2005; Jones et al. 2010b; Genzel et al. 2011; Swinbank et al. 2012; Fisher et al. 2014). The size and brightness of the star forming clumps in the spiral arm of *A1689B11* are comparable to those of $z > 1$ clumpy SF galaxies (Yuan, in preparation), producing a much clumpier appearance than local spiral arms. In addition, *A1689B11* is gravitationally lensed; the reconstructed source-plane morphology is much sharper than a non-lensed case, making it look different from most $z \sim 2$ SF galaxies.

To demonstrate the effect of large star forming clumps and observational effects such as surface brightness dimming and gravitational lensing on the appearance of spiral arms at high redshift, we manually redshift a local spiral galaxy to $z = 2.54$ and compare its redshifted morphology with the lensed and non-lensed case of *A1689B11* (Figure 9). The local spiral galaxy template (*G04-1*) is chosen from the DYNAMO (DYNAMics of Newly-Assembled Massive Objects) sample. The DYNAMO sample is a local analog of turbulent, clumpy disk galaxies at high redshift (Green et al. 2010, 2014; Bassett et al. 2014). *G04-1* is one of the few galaxies in the DYNAMO sample that have a spiral morphology. *G04-1* has similar half-light radius ($R_{1/2,H\alpha} = 2.7$ kpc), star forming clump size and brightness as *A1689B11* (Fisher et al. 2017). We use the HST $H\alpha$ narrow-band image (pixel size $0''.05$) of *G04-1* as it best represents the clumpy SF morphology (Figure 9, panel (1)). Figure 9-panel (2) illustrates the mocked HST/WFC3 (pixel size $0''.1$) morphology of *G04-1* at $z = 2.54$. We assume no intrinsic size evolution because the DYNAMO sample consists of compact objects in the local universe, anal-

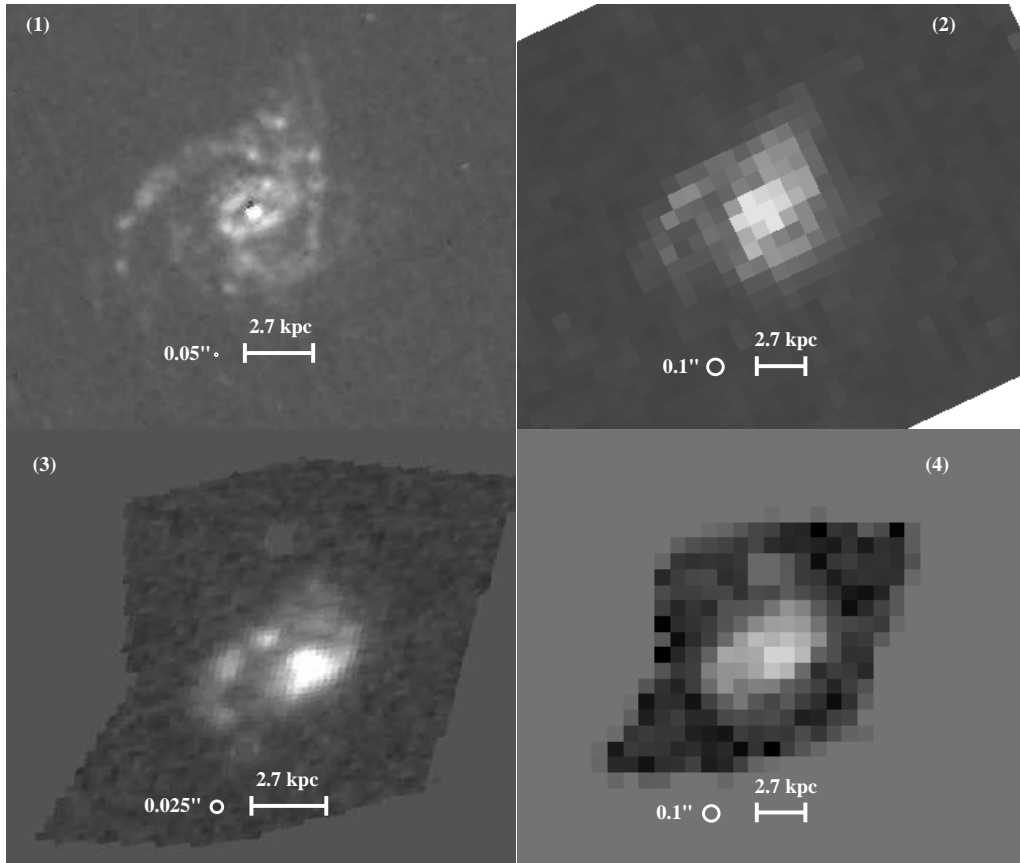


FIG. 9.— An illustration of the effect of large SF clumps, surface brightness dimming and gravitational lensing. (1) The HST/ACS narrow-band ($H\alpha$) image of a local spiral galaxy “ $G04-1$ ” that has large SF clumps similar to a $z \sim 2$ clumpy disks (from the DYNAMO sample). “ $G04-1$ ” has a half-light radius of 2.7 kpc, similar to $A1689B11$. (2) The mocked HST/WFC3 IR image of “ $G04-1$ ” after being redshifted to $z = 2.54$ without the effect of lensing. (3) The lens reconstructed morphology of $A1689B11$. (4) The morphology of $A1689B11$ as it would appear in HST/WFC3 IR band without the lensing magnification. The angular resolution and half-light radius are marked on each panel.

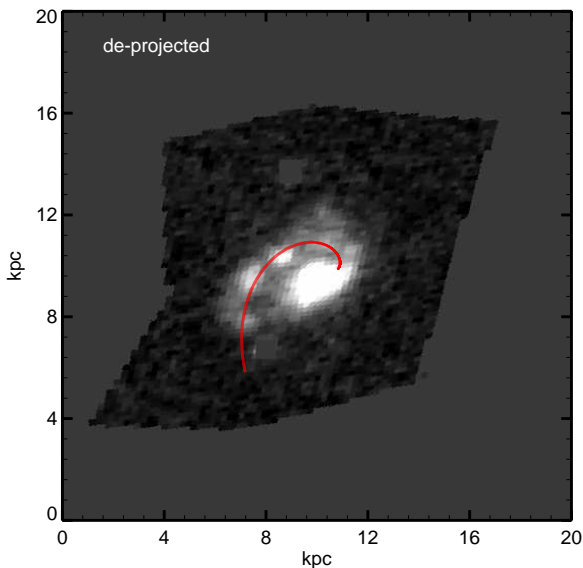


FIG. 10.— The best-fit logarithmic spiral function (red line) to the de-projected image of $A1689B11$. The best-fit pitch angle is $\theta = 37^\circ \pm 2^\circ$.

ogous to the sizes of galaxies $z \sim 2$. Because $G04-1$ is nearly face-on, the morphology of $A1689B11$ is de-projected in panels (3) and (4) for a better compari-

son. Panel (3) of Figure 9 shows the best source-plane reconstructed morphology of $A1689B11$ based on the HST/F814W band image. Owing to gravitational lensing, the effective spatial resolution on the source-plane is increased by a factor of 2-3 and the SNR of the image increased by ~ 7 . To compare with the non-lensed image of $G04-1$ at $z = 2.54$, we show in panel (4) what $A1689B11$ would have looked like without lensing magnification in the HST/WFC3 IR band.

Panel (1) and (2) of Figure 9 underline the importance of separating observational effects from intrinsic evolution of spiral arms. In the local universe, spiral galaxies are divided into three classes based on the number of arms: grand-design (two-arm), many-arm (multiple distinct global arms), and flocculent (multiple less distinct arms) (e.g., Elmegreen et al. 1982; Hart et al. 2017). According to the local image of Figure 9-panel (1), $G04-1$ is classified as a many-arm spiral, as are more than half of the local spiral galaxy population (Davis et al. 2014). As a result of surface brightness dimming, $G04-1$ would be mostly likely classified as a rare one-arm spiral that are only seen in 14% of local spirals (Davis et al. 2014). It is possible that $A1689B11$ has more arms like $G04-1$, but only the longest arm is visible at this redshift and with this magnification.

Panel (3) and (4) of Figure 9 demonstrate the power of gravitational lensing in bringing the otherwise unseen

spiral structures into focus. The spiral structure is still detectable in the redshifted image of $G04-1$, whereas the spiral arm of $A1689B11$ is barely distinguishable in the un-lensed case of panel (4). Note that although $G04-1$ and $A1689B11$ share similar galaxy size and SF clump brightness, $G04-1$ is ~ 10 times more massive than $A1689B11$ and has a large $H\alpha$ velocity dispersion of $\sim 50 \text{ km s}^{-1}$. There is noticeable differences in the morphological appearance of panel (3) and (4), implying that the intrinsic spiral structures of $A1689B11$ are less well-developed than $G04-1$. While a detailed comparison of the host properties of $G04-1$ and $A1689B11$ is beyond of the focus of the current paper, Figure 9 simply demonstrates that the morphology of $A1689B11$ is consistent with a spiral galaxy. The effect of gravitational lensing, surface brightness dimming and larger SF regions at high-redshift combined together to make the somewhat unique appearance of $A1689B11$.

Finally, separating spiral arms from clumpy irregular morphologies requires a quantitative classification scheme. Such a scheme is not available yet at high redshift. For example, there are two geometric parameters that are commonly used in quantifying local spiral arms: the number of arms (or harmonic modes) and the pitch angle. A large fraction of local spirals can be modeled by superpositions of logarithmic spiral functions. Based on this, automatic spiral arm finding and logarithmic function fitting tools have been developed and applied in local galaxy surveys (e.g., Davis et al. 2012; Davis & Hayes 2014; Shields et al. 2015; Hart et al. 2017). Our on-going effort of testing and adapting these tools at high-redshift will help to quantify spiral arms and to separate spiral from irregular structures objectively. We show in Figure 10 a preliminary best-fit logarithmic spiral function created by manually masking pixels that form the spiral arm. We find a large pitch angle ($\theta = 37^\circ \pm 2^\circ$) for the arm. We test the automated arm detection and fitting routine of SpArcFiRe (Davis & Hayes 2014) on $A1689B11$ and notice that the number of arms and pitch angle depend sensitively on the bulge-disk decomposition, lensing PSF reconstruction and masking of noise. We will report the pitch angle analysis in our future work (Yuan, in preparation). To conclude, we favor the interpretation of a spiral galaxy with primitive spiral arms developing in $A1689B11$ instead of a merger or an irregular galaxy.

4.3. Comparison with other samples

The total SFR ($22 M_\odot \text{ yr}^{-1}$) and stellar mass ($10^{9.8} M_\odot$) place $A1689B11$ as a typical SF main-sequence galaxy at $z \sim 2$. However, the kinematic properties of $A1689B11$ are quite “mature” compared to a $z > 1$ SF galaxies and are more akin to local spiral galaxies.

We compare the gas velocity dispersion and stellar mass of $A1689B11$ with other samples in Figure 11. We collect $H\alpha$ gas velocity dispersion measurements on disk galaxies from a few representative IFS surveys at $z \gtrsim 2$: the SINS disks (Förster Schreiber et al. 2006, 2009; Cresci et al. 2009), the KMOS-3D disks (Wisnioski et al. 2015), and gravitationally lensed galaxies (Jones et al. 2010b; Livermore et al. 2015). Only the rotation-dominated ($V_{rot}/V_\sigma > 1$) disks from these surveys are shown. The $z \sim 0$ sample is taken from a local reference sample of rotating spiral galaxies with $H\alpha$ velocity dispersion mea-

surements (Epinat et al. 2010).

The local spirals have low $H\alpha$ velocity dispersions irrespective of their stellar masses. There is a large spread in velocity dispersions of $z \gtrsim 2$ disks, with a median value that is ~ 2.5 times larger than local spirals. We also mark the thermal broadening of $H\alpha$ emitting gas (10^4 K) as the minimal velocity dispersion that can be measured from the $H\alpha$ line. The systematic uncertainty in comparing velocity dispersions across samples could also contribute to the large scatter in high-redshift samples. The major systematics come from the method used to correct for beam-smearing, the disk model, and the radius at which the velocity dispersion is measured and the weighting used. The velocity dispersion of $A1689B11$ is considerably lower than lensed galaxies of similar masses. It is also much lower than the median of all $z \gtrsim 2$ samples. The large V_{rot}/V_σ ($\sim 9-13$) of $A1689B11$ is similar to local spirals and $\sim 2-5$ times larger than the median value of SF galaxies at $z \gtrsim 1$ (e.g., Cresci et al. 2009; Swinbank et al. 2017).

One of the major observational results from IFS surveys of high-redshift ($1 \lesssim z \lesssim 3$) galaxies is that the intrinsic gas velocity dispersions as usually measured from optical ionized gas are significantly higher than local SF galaxies (e.g., Law et al. 2009; Lehnert et al. 2009; Förster Schreiber et al. 2009; Genzel et al. 2011; Wisnioski et al. 2015). Popular explanations for the enhanced velocity dispersions include star formation feedback (Lehnert et al. 2009; Green et al. 2014), gravitational instability (Agertz et al. 2009; Ceverino et al. 2010), a combination of star formation feedback driven and gravitationally driven turbulence (Krumholz & Burkhardt 2016; Krumholz et al. 2017), a multitude of physical drivers of turbulence (Federrath & Klessen 2012; Federrath et al. 2017a) and cosmological cold gas accretion (Dekel et al. 2009; Bournaud & Elmegreen 2009; Genzel et al. 2011; Genel et al. 2012).

If the ionized gas turbulence of $A1689B11$ is directly driven by energy ejected from star formation through supernova, then its SFR surface density suggests a velocity dispersion of $V_\sigma \sim 40-55 \text{ km s}^{-1}$ (e.g., Lehnert et al. 2009; Green et al. 2014), depending on how the energy is dissipated into the ISM and assuming a conservative supernova feedback efficiency of 0.25 (Dib et al. 2006; Zhou et al. 2017). In this simple model, the small velocity dispersion of $A1689B11$ would imply that the supernova feedback efficiency is a factor of ~ 2 lower than local galaxies, which is difficult to explain.

In the slightly more complicated feedback driven turbulence model of Krumholz & Burkhardt (2016), where V_σ is a function of SFR and Q_{gas} , the small velocity dispersion of $A1689B11$ can be reproduced at a $Q_{gas} \sim 0.5$, matching our rough estimation of Q_{gas} in Section 3.3. The gravity driven turbulence model of Krumholz & Burkhardt (2016) is disfavored because the small V_σ would require an unrealistically large gas fraction ($50\% < f_{gas} \sim 100\%$). Future direct observations of molecular gas and dust maps of $A1689B11$ would provide a more robust measurement on Q_{gas} and gas fraction to distinguish various turbulence driven models.

We also show in Figure 11 the location of the spiral galaxy $BX442$ at $z=2.17$ from Law et al. (2012). $BX442$ has a high velocity dispersion ($V_\sigma \sim 70 \text{ km s}^{-1}$) and is

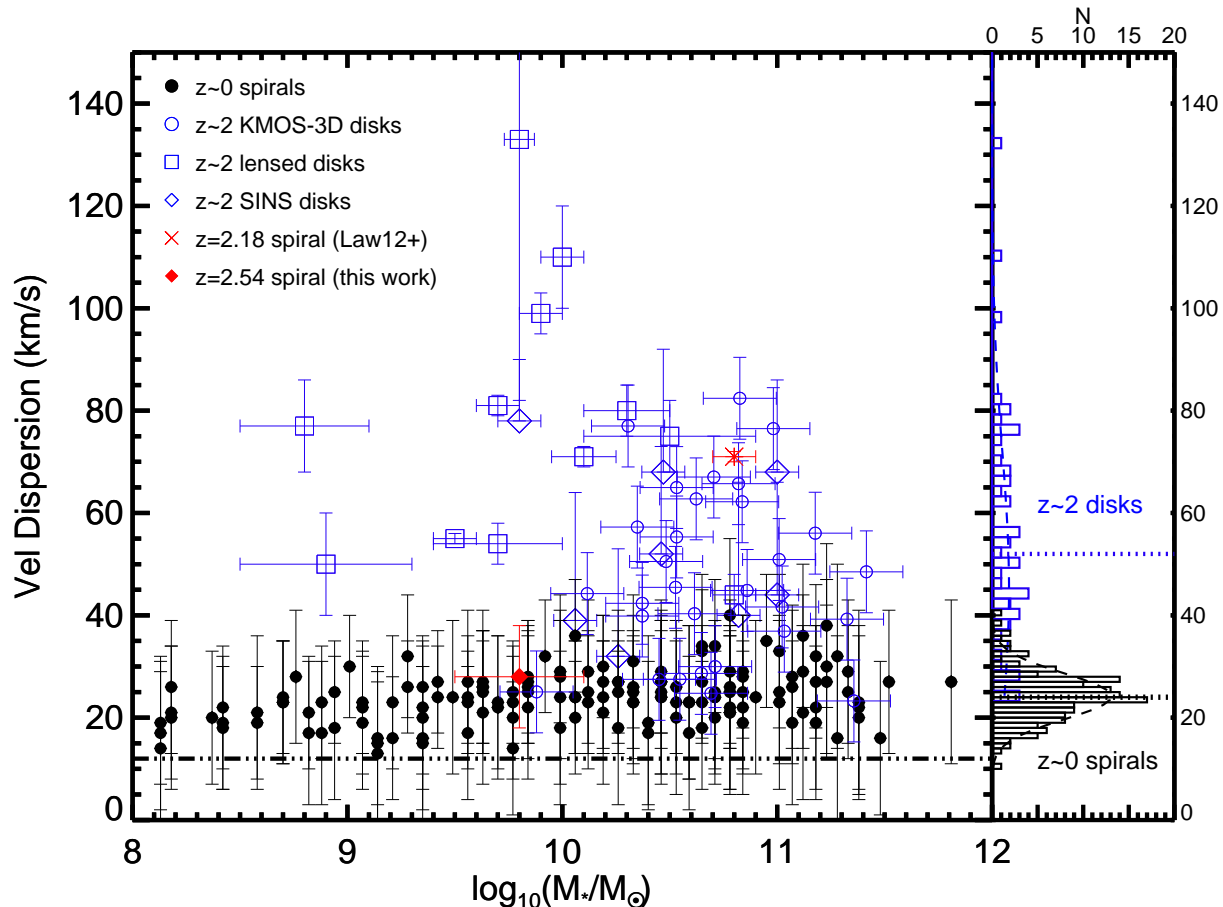


FIG. 11.— $H\alpha$ velocity dispersion versus stellar mass for *A1689B11* (the red diamond) and comparison samples at $z \sim 0$ and $z > 2$. The blue empty diamonds are the $z > 2$ SINS disks (Förster Schreiber et al. 2006, 2009; Cresci et al. 2009); the blue empty circles are the $z > 2$ KMOS-3D disks (Wisnioski et al. 2015), the blue empty squares are gravitationally lensed galaxies at $z > 2$ (Jones et al. 2010b; Livermore et al. 2015). The local spiral galaxy sample (black filled circles) is from Epinat et al. (2010). The red cross shows the only other known $z > 2$ spiral, BX442, detailed in Law et al. (2012). *A1689B11* is marked by the red filled diamond. The horizontal dash-dotted line indicates the minimal velocity dispersion that can be probed by $H\alpha$ lines because of the thermal broadening ionized gas (10^4 K). The panel on the right shows the velocity dispersion distribution of the $z \sim 0$ and $z \sim 2$ samples.

thought to have a short-lived spiral triggered by a minor merger (Law et al. 2012). Two conditions must be satisfied according to the minor-merger trigger mechanism. First the galaxy must be massive enough to stabilize the formation of an extended disk. Second, a nearby merging satellite must be properly orientated and sufficiently massive to excite the spiral feature. *A1689B11* is ~ 10 times less massive than *BX442* and is not consistent with a merging system as discussed in Section 4.1. Even if the minor merger scenario works for *BX442*, an alternative mechanism is required to account for the existence of *A1689B11*.

4.4. Spiral Arm and Thin Disk formation

Being an outlier in both the velocity dispersion distribution and morphology (spiral structure) of high-redshift galaxies, *A1689B11* provides a few interesting angles to revisit the formation of spiral arms.

Classic theories of spiral arm formation require a dynamically cool and thin disk. In the paradigm of the density wave theory, the amplitude of the induced density wave becomes too weak if the gas velocity dispersion is too high or if the disk is too thick (e.g., Lin &

Shu 1964; Bertin & Lin 1996; Elmegreen & Thomasson 1993; Rafikov 2001; Bottema 2003; Sellwood 2014). In the alternative theory of swing amplification or self-gravitational instability, a large velocity dispersion (in the context of a large Toomre- Q parameter) dampens the amplification and a thick disk reduces the disturbance gravity that are seeds of the swing amplification (e.g., Julian & Toomre 1966; Elmegreen & Thomasson 1993; Bottema 2003; Sellwood 2014). In the local universe, almost all spiral arms reside in the thin disk (scale height 200-300 pc) with a dynamically cool stellar ($V_\sigma(\text{star}) \sim 20$ km/s) and gas component ($V_\sigma(H\alpha) \sim 20 - 25$ km/s) (Glazebrook 2013). For stellar disks that are dynamically “warm” and have a non-negligible thickness, spiral activity develops mostly in the cold gas component (Bertin & Lin 1996). The exceptions are mergers or bar induced spirals, where the external driver plays a more important role than internal disk dynamics (e.g., Kormendy & Norman 1979; Bottema 2003; Dobbs et al. 2010).

Both the density wave and swing amplification theory have seen a certain degree of success in explaining local spirals where the model assumptions are easily satisfied.

The logarithmic spiral function predicted from the density wave theory has been widely used and confirmed in observations of local spirals, whereas swing amplification is commonly used to explain transient and less regular spirals. The main difference is that density wave theory predicts quasi-steady long-lived arms (wave arm) and swing amplification or its various form of gravitational instability predicts short-lived reoccurring arms (material arm). It is difficult to constrain the lifetime of the spiral arm observationally, thus hard to distinguish the two theories (e.g., Sellwood 2011). Current simulations seem to prefer either swing amplification or merger in explaining the formation of spiral arms at high redshift (e.g., Law et al. 2012; Fiacconi et al. 2015). For example, self-gravitational instability and fragmentation has been proposed to model the local DYNAMO spiral G04 – 1 (Section 4.2) that has a high velocity dispersion and clumpy SF regions analogous to high-redshift disks (Inoue & Yoshida 2017). Minor merger is thought to trigger the short-lived spiral structure of *BX442* (Law et al. 2012).

It is reasonable to expect that the quasi-steady spiral arm from the density wave theory is suppressed at high redshift because of the high gas velocity dispersions and geometrically thick disks (Cresci et al. 2009; Elmegreen & Elmegreen 2006). The rarity of spirals at $z \gtrsim 2$ could be partly accounted for by the short-lived arms from either gravitational instability or mergers. On the other hand, the discovery of a dynamically cool disk like *A1689B11* could mean that the condition for the classic density wave spirals to develop can exist at $z > 2$. The small velocity dispersion of *A1689B11* is indicative of a thin disk (Figure 11). Our preliminary fitting of a logarithmic spiral function (Figure 10) is consistent with a density wave triggered primitive spiral arm of *A1689B11*. We speculate that *A1689B11* belongs to a population of rare spiral galaxies at $z \gtrsim 2$ that mark the earliest epoch of thin disk formation (Freeman & Bland-Hawthorn 2002; Kraljic et al. 2012; Freeman 2012; Elmegreen et al. 2017). Note that the spiral structure of *A1689B11* would not be visible with current observational capacity without gravitational lensing (Figure 9). Future observations with the James Webb Space Telescope (JWST) will help to reveal this population and investigate the earliest onset of spiral arms and thin disks.

4.5. The origin of spiral arms in cosmological zoom-in simulations

Modern theoretical efforts in understanding the origin of spiral arms have focused on semi-analytic or N-body simulations in isolated disks (e.g., Wada et al. 2011; Baba et al. 2013; D’Onghia et al. 2013; D’Onghia 2015). Cosmological simulations do not have sufficient resolutions to trace the detail dynamics of spiral arms, however, it is interesting to explore the spiral arm formation from the context of cosmological simulations where the role of environment is included.

Recent cosmological zoom-in simulations (Cen 2014) provide valuable insights into the emergence of the Hubble sequence across cosmic time. Cen (2014) shows the rich physics of cold gas accretion dynamics onto galaxies and evolutionary trends. At $z \gtrsim 2$, the average in-situ cold gas accretion streams through the galactic halo virial sphere surface can be characterized by: (1) multiple cold

streams, (2) high accretion rates, (3) low angular momenta, and (4) high gas densities. Among these four parameters, it is suggested that spiral structures are most sensitive to (1) the number of concurrent streams in the gas accretion. Flocculent spirals only begin to significantly appear at $z \sim 1 - 2$ when the number of major gas streams are about two to three and two-arm grand-design spiral galaxies appear at $z \leq 1$ when the number of major cold streams reduces to one. In the framework of Cen (2014), the average number of gas streams decreases from high to low redshift. However, at each redshift, there is a distribution of in-situ environments pertaining to the range of cold streams. It is therefore possible to find “evolved” spiral galaxies at high- z and “unevolved” high-redshift analog galaxies at low- z at the tail of the environmental distributions.

In the framework of Cen (2014), spiral *A1689B11* may be formed in an in-situ environment that is characterized by probably one major stream with high accretion rates and gas densities for an extended period of time. Because the number of significant cold streams is correlated with the degree of interactions among galaxies, it is expected that the velocity dispersion of the stellar disk would be positively correlated with the number of cold streams (Cen 2014). This scenario is consistent with accretion energy being one of the main drivers of the turbulence in disks at $z \sim 2$ (e.g., Birnboim & Dekel 2003; Dekel et al. 2009; Bournaud & Elmegreen 2009; Genel et al. 2012).

However, galaxy growth through cold-mode accretion has been challenged by other cosmological simulations (e.g., Nelson et al. 2013; Genel et al. 2014). Whether the Hubble sequence originates from the cold flow process or other mechanisms such as mergers and feedback is highly controversial (e.g., Genel et al. 2015). Moreover, extra care must be given when studying the origin of the spiral arms in cosmological simulations. Spiral features in current cosmological simulations may originate from unphysical perturbations and are sensitive to the resolution and the detailed prescriptions of ISM models. With a larger sample of high-redshift spirals in observations and larger volumes of cosmological zoom-in simulations with sub-pc resolutions, we should be able to answer the question of whether the in-situ environment plays a critical role in the formation of spiral galaxies.

5. SUMMARY AND FUTURE WORK

We report NIFS/Gemini observations on a $z = 2.54$ gravitationally lensed spiral galaxy *A1689B11*. It is the highest redshift spiral galaxy observed with the highest spatial resolution and spectroscopic depth to date. *A1689B11* shows a primitive spiral arm that is scarcely seen in other galaxies at $z \gtrsim 2$. Regarding the SFR, size and stellar mass, *A1689B11* is representative of a $z \sim 2$ SF galaxy. In contrast, the $H\alpha$ kinematic field shows striking similarities to $z \sim 0$ isolated late-type spiral galaxies. It shows an ordered rotation ($V_c = 200 \pm 12$ km s⁻¹) and uniformly small velocity dispersions ($V_{\sigma, \text{mean}} = 23 \pm 4$ km s⁻¹ and $V_{\sigma, \text{outer-disk}} = 15 \pm 2$ km s⁻¹). The low gas velocity dispersion is consistent with the classic density wave theory that spiral arms form in dynamically cold and thin disks. We speculate that *A1689B11* belongs to a population of rare spiral galaxies at $z \gtrsim 2$ that mark the formation epoch of thin disks. Future ob-

TABLE 1
PHYSICAL PROPERTIES OF *A1689B11*

RA (J2000)	13:11:33.336
DEC(J2000)	-01:21:06.9
<i>Spectroscopic data measurements:</i>	
Redshift ($z_{H\alpha}$)	2.540
Velocity dispersion (V_{σ} , mean)	$23 \pm 4 \text{ km s}^{-1}$
Velocity dispersion (V_{σ} , outer disk 1-2 kpc)	$15 \pm 2 \text{ km s}^{-1}$
Star formation rate ($\text{SFR}_{H\alpha}$)	$22 \pm 2 M_{\odot} \text{ yr}^{-1}$
Dust attenuation ($E(B-V)_{\text{Balmer decrement}}$)	0.73
<i>GALFIT/exponential disk best-fit:</i>	
Inclination (i)	55 ± 10 degrees
Position angle (PA)	-36 ± 6 degrees
Scale length (r_s)	1.3 ± 0.4 kpc
Half-light radius ($R_{1/2}$)	2.6 ± 0.7 kpc
Effective radius (R_e)	2.0 ± 0.4 kpc
<i>2D velocity disk model best-fit:</i>	
Inclination (i , velocity)	51 ± 2 degrees
Position angle (PA, velocity)	-37 ± 2 degrees
Radius rotation (R_t , velocity)	1.7 ± 0.1 kpc
Rotation Velocity (V_{rot})	$200 \pm 12 \text{ km s}^{-1}$
<i>Inferred physical parameters:</i>	
Star formation rate (SFR_{SED})	$22 \pm 3 M_{\odot} \text{ yr}^{-1}$
Stellar mass (M_{star})	$10^{9.8 \pm 0.3} M_{\odot}$
Dynamical mass (M_{dyn})	$10^{10.2} M_{\odot}$
Average SFR surface density (Σ_{SFR})	$0.3 M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$
Average gas surface density (Σ_{gas})	$158 M_{\odot} \text{ pc}^{-2}$
Gas fraction (f_{gas})	~ 0.18
Toomre parameter Q_{gas}	~ 0.7

NOTES. — All values have been corrected for lensing magnifications. PA are defined as: PA=0 when the major axis is up and positive if rotated counterclockwise. Assumptions for inferred values are described in Section 3.

servations with JWST will help to reveal this population and investigate the earliest onset of spiral arms.

Our follow-up work include a detailed study on the angular momentum and SF clump properties by including our recent OSIRIS observation on the more magnified image B11.2 of *A1689B11*. We are exploring a more robust method of bulge/disk decomposition of *A1689B11* and plan to investigate the bulge to pitch angle correlation in a follow-up paper (Yuan, in prep). Our on-going

IFS observations on a larger sample of non-lensed spiral galaxies (~ 30) at $z \gtrsim 2$ that we recently discovered will help to validate/reject our speculations about the origin and number density of spiral galaxies like *A1689B11*.

We thank the anonymous referee for his/her report, which helped to restructure and improve the work significantly. This work is a tribute to the late Peter McGregor who built NIFS/Gemini and taught TY how to reduce the NIFS data. This research was conducted by the Australian Research Council Centre of Excellence for All Sky Astrophysics in 3 Dimensions (ASTRO 3D), through project number CE170100013. TY thanks useful discussions with Ken Freeman, Lars Hernquist, Enrico Teodoro, Yusuke Fujimoto, Ben Davis and the GEARS3D group. TY acknowledges the support from the ASTRO 3D fellowship. BG gratefully acknowledges the support of the Australian Research Council as the recipient of a Future Fellowship (FT140101202). CF acknowledges funding provided by the ARC Discovery Projects (grants DP150104329 and DP170100603). JR acknowledges support from the ERC starting grant 336736-CALENDS. LK gratefully acknowledges support from an Australian Research Council (ARC) Laureate Fellowship (FL150100113). RC acknowledges grants NNX12AF91G and AST15-15389. YB acknowledges ISF grant 1059/14. We thank Stuart Ryder for his assistance with the Gemini observations.

The data for this work is based on observations obtained at the Gemini Observatory, which is operated by the Association of Universities for Research in Astronomy, Inc., under a cooperative agreement with the NSF on behalf of the Gemini partnership: the National Science Foundation (United States), the National Research Council (Canada), CONICYT (Chile), Ministerio de Ciencia, Tecnología e Innovación Productiva (Argentina), and Ministério da Ciência, Tecnologia e Inovação (Brazil). The authors wish to recognize and acknowledge the very significant cultural role and reverence that the summit of Mauna Kea has always had within the indigenous Hawaiian community.

Facilities: Gemini (NIFS).

APPENDIX

We show the source-plane reconstructed multi-wavelength morphology from image B11.1 and B11.2 in Figure 12 and Figure 13.

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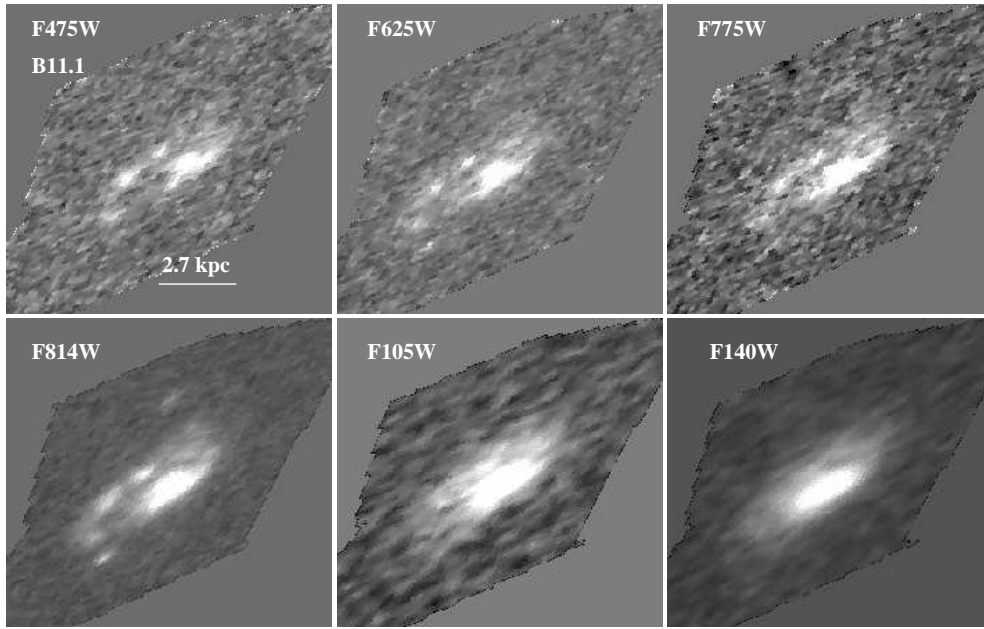


FIG. 12.— Multi-band HST images of B11.1 on the source-plane.

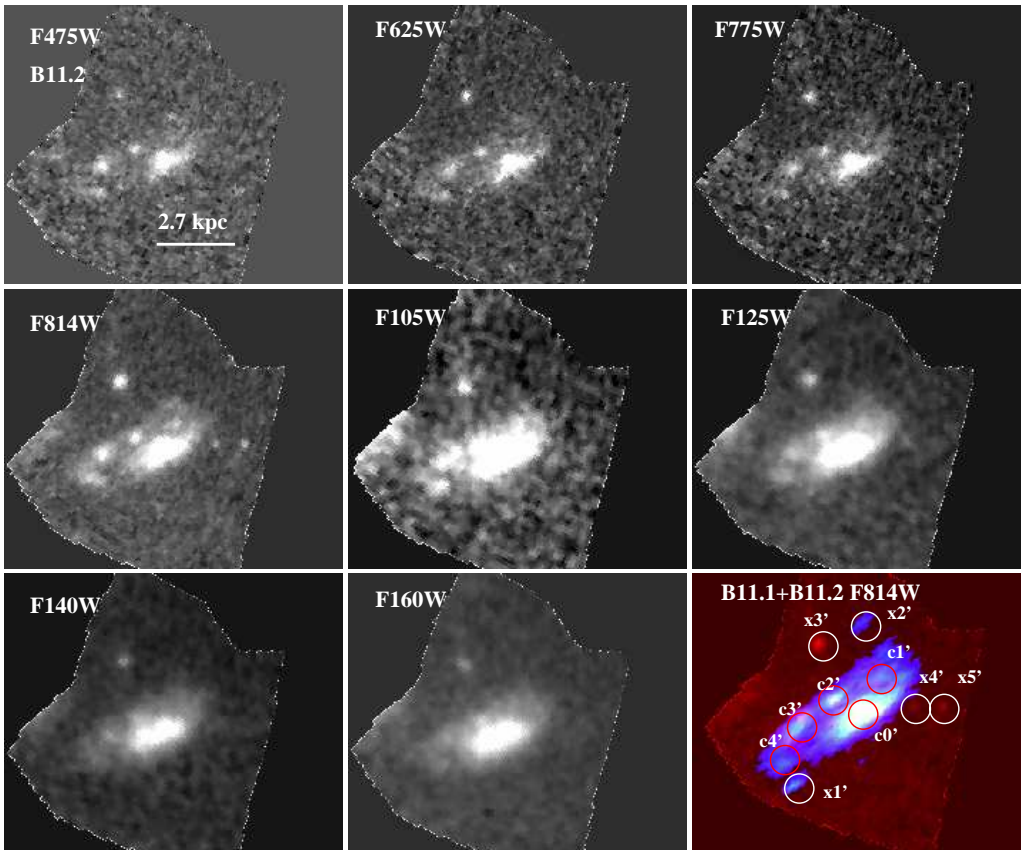


FIG. 13.— Multi-band HST images of B11.2 on the source-plane. The last panel combines the F814W-band morphology of B11.1 and B11.2. The clumps that are associated with the spiral host are marked in red circles whereas the interlopers are marked in white circles (see also Section 3.2 and Figure 7).

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